



External Inverse-Compton Emission Associated with Extended and Plateau Emission of Short Gamma-Ray Bursts: Application to GRB 160821B

B. Theodore Zhang^{1,2,3}, Kohta Murase^{1,2,3,4}, Chengchao Yuan^{1,2,3}, Shigeo S. Kimura^{5,6}, and Peter Mészáros^{1,2,3}

¹Department of Physics, Pennsylvania State University, University Park, PA 16802, USA

²Department of Astronomy & Astrophysics, Pennsylvania State University, University Park, PA 16802, USA

³Center for Multimessenger Astrophysics, Institute for Gravitation and the Cosmos, Pennsylvania State University, University Park, PA 16802, USA

⁴Center for Gravitational Physics, Yukawa Institute for Theoretical Physics, Kyoto University, Kyoto, Kyoto 606-8502, Japan

⁵Frontier Research Institute for Interdisciplinary Sciences, Tohoku University, Sendai 980-8578, Japan

⁶Astronomical Institute, Tohoku University, Sendai 980-8578, Japan

Received 2020 December 18; revised 2021 January 23; accepted 2021 January 27; published 2021 February 22

Abstract

The recent detection of TeV photons from two gamma-ray bursts (GRBs), GRB 190114C and GRB 180720B, has opened a new window for multimessenger and multiwavelength astrophysics of high-energy transients. We study the origin of very high energy (VHE) γ -rays from the short GRB 160821B, for which the MAGIC Collaboration reported a $\sim 3\sigma$ statistical significance. Short GRBs are often accompanied by extended and plateau emission, which is attributed to internal dissipation resulting from activities of a long-lasting central engine, and Murase et al. (2018) recently suggested the external inverse-Compton (EIC) scenario for VHE counterparts of short GRBs and neutron star mergers. Applying this scenario to GRB 160821B, we show that the EIC flux can reach $\sim 10^{-12}$ erg cm $^{-2}$ s $^{-1}$ within a time period of $\sim 10^3$ – 10^4 s, which is consistent with the MAGIC observations. EIC γ -rays expected during the extended and plateau emission will be detectable with greater significance by future detectors such as the Cerenkov Telescope Array. The resulting light curve has a distinguishable feature, where the VHE emission is predicted to reach the peak around the end of the seed photon emission.

Unified Astronomy Thesaurus concepts: [Gamma-ray bursts \(629\)](#); [Gamma-ray sources \(633\)](#); [Gamma-ray astronomy \(628\)](#); [High energy astrophysics \(739\)](#); [Relativistic jets \(1390\)](#); [Gravitational wave astronomy \(675\)](#); [Non-thermal radiation sources \(1119\)](#)

1. Introduction

Very high energy (VHE) γ -rays (with energies higher than ~ 0.1 TeV) represent the most energetic part of the currently observed electromagnetic (EM) spectrum. The VHE γ -rays play an important role in high-energy multimessenger and multiwavelength astrophysics (e.g., Inoue et al. 2013; Mészáros et al. 2019; Murase & Bartos 2019; Hinton & Ruiz-Velasco 2020). Imaging Atmospheric Cerenkov Telescopes (IACTs) detect Cerenkov light that is produced during the development of extensive air showers as the VHE γ -rays enter Earth's atmosphere (e.g., Hinton 2009; Lorenz et al. 2012). The detection of VHE γ -rays from GRB 190114C by the Major Atmospheric Gamma Imaging Cerenkov (MAGIC) telescopes (Acciari et al. 2019a, 2019b) and GRB 180720B by the High Energy Stereoscopic System (H.E.S.S.; Abdalla et al. 2019) has opened a new window for the exploration of the physics of relativistic shocks involving particle acceleration (Mészáros 2006; Kumar & Zhang 2014).

On the other hand, the discovery of the first double neutron star (NS) merger event GW170817, associated with GRB 170817A (Abbott et al. 2017a), is a milestone in the multimessenger astronomy, which was initially detected through gravitational waves (GWs; Abbott et al. 2017b) and later observed through EM emission from radio to γ -rays (Abbott et al. 2017c). VHE γ -ray emission has also been recently discussed in this context (Murase et al. 2018; Kimura et al. 2019), and the GW follow-up observations by Cerenkov telescopes, e.g., the Cerenkov Telescope Array (CTA), are promising in the near future (Kakuwa et al. 2012; Inoue et al. 2013; Bartos et al. 2019). In the synchrotron self-Compton

(SSC) scenario, the same population of electrons that emit synchrotron photons in the dissipation region of the relativistic outflow upscatter these photons to much higher energies by a factor of $\sim \gamma_e^2$, where γ_e is the electron Lorentz factor (e.g., Mészáros & Rees 1994; Sari & Esin 2001; Wang et al. 2001; Zhang & Mészáros 2001; Acciari et al. 2019b; Abdalla et al. 2019). In the presence of long-lasting central engine activities, late-prompt photons that are related to the extended and/or plateau emission can be upscattered to the VHE band by high-energy electrons accelerated at the external forward shock via external inverse-Compton (EIC) emission (Murase et al. 2018, for NS mergers). On the other hand, if the prolonged jets dissipate inside a cocoon, VHE γ -rays via upscattering of thermal cocoon photons in the jet have also been predicted (Kimura et al. 2019).

The MAGIC experiment utilizes two IACTs to detect VHE γ -rays within the energy range from ~ 50 GeV to ~ 50 TeV (Aleksić et al. 2016a, 2016b). Besides the confirmed detection of GRB 190114C, there are three short gamma-ray bursts (GRBs), GRB 061217, GRB 100816A, and GRB 160821B, with redshifts $z < 1$ and time delays after the GRB trigger < 1 hr that have been followed up by the MAGIC telescopes under adequate conditions (Acciari et al. 2019a). In particular, the MAGIC Collaboration recently reported an excess of VHE γ -rays from the direction of the short GRB 160821B within the time window of 24 s to 4 hr after the trigger (Inoue et al. 2019; Acciari et al. 2020), and there was a $\sim 3\sigma$ signal at $t \sim 10^4$ s. GRB 160821B is one of the nearest short GRBs identified by the Neil Gehrels Swift Observatory (Levan et al. 2016), which is located at the outskirts of the host galaxy with a measured

redshift of $z \sim 0.16$ (or a luminosity distance of ~ 780 Mpc). Although the detection is still tentative, GRB 160821B may be the first short GRB detected in the VHE band.

The γ -ray and X-ray light curves observed by Swift often show a period of extended emission in the early X-ray afterglow phase with a duration of ~ 300 s (Norris & Bonnell 2006; Berger 2014). The existence of the extended emission component supports the argument that the central engine activity lasts for a longer time than the prompt emission, which could be explained by a magnetar (Dai et al. 2006) or black hole accretion (Kisaka & Ioka 2015). Short GRBs with extended emission usually show a plateau emission phase at later times (Gompertz et al. 2013, 2014; Kisaka et al. 2017). The X-ray flux during the plateau phase can be explained within the external forward shock model, considering a complicated structured jet with an appropriate initial Lorentz factor (Lamb et al. 2019; Troja et al. 2019, for GRB 160821B), or else through a refreshed shock scenario in which the forward shock is continuously replenished via the collision with the slower but more energetic portions of the ejecta (Matsumoto et al. 2020). Alternatively, both extended and plateau emission components are attributed to late-prompt emission from the long-lasting internal dissipation (Ghisellini et al. 2007; Murase et al. 2011; Kisaka & Ioka 2015; Kisaka et al. 2017), where the external shock emission can be outshined by the late-prompt emission.

In this letter, we explore the origin of VHE γ -rays from GRB 160821B considering both of the SSC and EIC scenarios. In Section 2, we show that the SSC scenario is disfavored to explain the observed VHE γ -rays given the multiwavelength constraints. In Section 3, we present the results of the EIC scenario considering extended and plateau emission as seed photons. We discuss the implications of our results and give a summary in Section 4.

2. VHE γ -Rays in the Synchrotron Self-Compton Scenario

First, we consider the standard external forward shock model (Mészáros & Rees 1997; Sari et al. 1998), where electrons are accelerated to higher energies via the diffusive shock acceleration mechanism. In the SSC scenario, synchrotron photons emitted from these electrons are upscattered via the inverse-Compton (IC) emission mechanism. Following Zhang et al. (2020), the dynamics of the outflow is calculated by solving the 1D differential equation. The steady-state electron energy spectrum is derived by solving the continuity equation, and the observed flux from relativistic outflows is calculated via the integration along the equal-arrival-time surface (EATS). See Zhang et al. (2020) for details.

The best-fit parameters utilizing the multiwavelength data from radio to X-rays suggest relatively larger values of $\epsilon_B \sim 10^{-2}$ (Lamb et al. 2019; Troja et al. 2019), where ϵ_B is the energy fraction of the internal energy that is converted into the magnetic energy. However, the resulting SSC flux turns out to be too low for $\epsilon_e/\epsilon_B \sim 10$. In this work, we adopt $\epsilon_e = 0.3$ and $\epsilon_B = 10^{-5}$ to achieve the SSC dominance (e.g., Abdalla et al. 2019; Derishev & Piran 2019). Note that such a small value of ϵ_B can be consistent with the values extrapolated from the value obtained by numerical simulations (Vanthieghem et al. 2020).

The relation between the peak flux of IC and seed photons is approximately written as $F_{\epsilon, \max}^{\text{IC}} \sim \tau_e f_e F_{\epsilon, \max}^{\text{seed}}$ in the Thomson regime, where $\tau_e \sim \sigma_T n_{\text{ex}} R$ is the Thomson optical depth in the shocked region, σ_T is the Thomson cross section, f_e is the

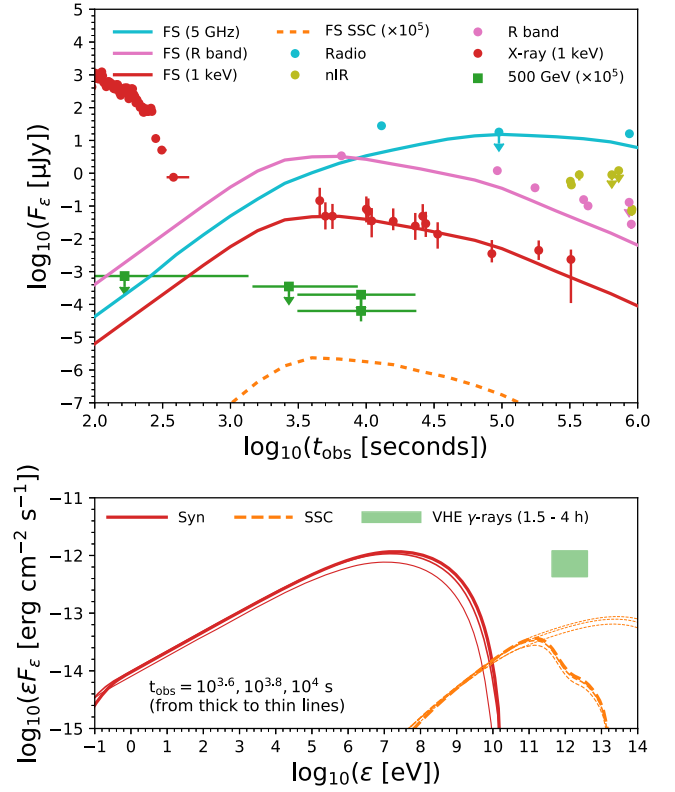


Figure 1. Top panel: multiwavelength light curves from radio, optical, X-ray, and VHE bands at 500 GeV in the SSC scenario. The radio and optical are taken from Lamb et al. (2019) and Troja et al. (2019), late-time X-ray data are taken from Acciari et al. (2020), while the early X-ray data at extended emission phase are taken from the public online repository assuming spectral index $\Gamma_X = 2.5$ (Evans et al. 2010). The green circles are the upper limit of the flux at 500 GeV assuming a power-law spectrum with spectral index $\Gamma_\gamma = 2$ (Acciari et al. 2020). Bottom panel: energy spectra at $t = 10^{3.6-10^4}$ s for different components, synchrotron emission, and SSC emission. The thick dashed lines take into account both the internal $\gamma\gamma$ absorption and EBL attenuation (Kneiske et al. 2004), while the thin lines ignore the EBL attenuation. The rectangle band shows the integrated flux level measured by MAGIC from 1.5 to 4 hr (Acciari et al. 2020). The physical parameters are $\mathcal{E}_k = 3 \times 10^{51}$ erg, $\Gamma_0 = 40$, $n_{\text{ex}} = 0.05 \text{ cm}^{-3}$, $\epsilon_e = 0.3$, $\epsilon_B = 10^{-5}$, $f_e = 0.5$, $s = 2.3$, and $\theta_j = 0.15$.

number fraction of electrons that can be accelerated, n_{ex} is the number density of external material, and R is the size of the emission region (Sari & Esin 2001; Wang et al. 2001; Zhang & Mészáros 2001). Thus, one expects that higher external density environments are preferred for strong IC emission. Unlike long GRBs that may occur in dense environments as a result of extensive wind losses, short GRBs usually occur at the outskirts of the host galaxy owing to the kick after the supernova explosions. Observations have revealed that the external matter density surrounding short GRBs usually has $n_{\text{ex}} \sim 10^{-3} - 1 \text{ cm}^{-3}$, depending on the merger timescale and kick velocity (Berger 2014), and the mean value is $n_{\text{ex}} \sim 3 \times 10^{-3} - 1.5 \times 10^{-2} \text{ cm}^{-3}$ (Fong et al. 2015). The best-fit parameters derived from the fitting to the multiwavelength data of GRB 160821B in the forward shock model show that the external matter density can be as low as $\sim 10^{-4} \text{ cm}^{-3}$ (Lamb et al. 2019; Troja et al. 2019) but can still be as high as $n_{\text{ex}} \sim 0.1 \text{ cm}^{-3}$ (Schroeder et al. 2020).

In Figure 1, we show the light curves (top panel) and energy spectrum (bottom panel) of the afterglow emission predicted in the SSC scenario. In the top panel, we show multiwavelength

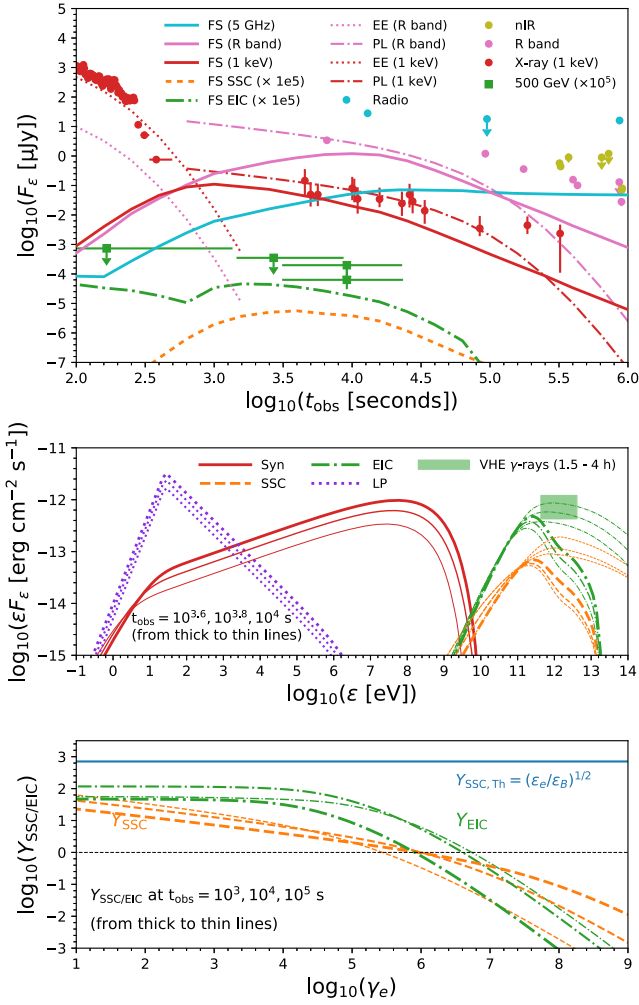


Figure 2. Top panel: multiwavelength light curves from radio, optical, X-ray, and VHE band at 500 GeV in the EIC scenario. The data shown are the same as in Figure 1. We include late-prompt emission consisting of extended and plateau emission fitted with Equations (1) and (2), and VHE γ -rays from the EIC process. Middle panel: same as Figure 1, but we show the broadband spectra for $t_{\text{obs}} = 10^{3.6}, 10^{3.8}, 10^4$ s. Bottom panel: Compton Y parameter, Y_{SSC} (dashed line) and Y_{EIC} (dashed-dotted line), as a function of electron Lorentz factor γ_e ; the blue solid line is the Y_{SSC} estimated in the Thomson regime. The physical parameters are $\mathcal{E}_k = 2 \times 10^{51}$ erg, $\Gamma_0 = 60$, $n_{\text{ex}} = 0.1 \text{ cm}^{-3}$, $\epsilon_e = 0.5$, $\epsilon_B = 10^{-6}$, $f_e = 0.1$, $s = 2.5$, and $\theta_j = 0.1$.

light curves from the external forward shock. The early radio data can be explained by the emission from reverse shock, and the late optical to infrared data may have contributions from a kilonova ejecta (Lamb et al. 2019; Troja et al. 2019). In the bottom panel, we show broadband energy spectra for synchrotron and SSC components at $t \sim 10^{3.4} - 10^4$ s. The thick γ -ray energy spectra take into account the effect of $\gamma\gamma$ absorption, where the extragalactic background light (EBL) attenuation is more important than the internal absorption by ambient photons. We also added the MAGIC light curve and the energy spectrum with the EBL correction for comparison (Inoue et al. 2019; Acciari et al. 2020).

We see, however, that from the SSC mechanism the predicted flux of VHE γ -rays is ~ 1 – 2 orders of magnitude lower than the MAGIC detection threshold within the time range from $t \sim 10^3$ s to $t \sim 10^4$ s. The Compton Y parameter, which is expressed as the ratio of the SSC power to the synchrotron power, can be estimated to be

$Y_{\text{SSC}}(\gamma_e) \propto U'_{\text{syn}}[\epsilon'_\gamma < \epsilon'_{\text{KN}}]/U'_B$, where primed quantities are measured in the jet comoving frame, U'_{syn} is the comoving synchrotron photon density, ϵ'_{KN} is the comoving Klein–Nishina break energy (Murase et al. 2011), and $U'_B = B^2/8\pi$ is the comoving magnetic energy. We find that the value of Y_{SSC} is lower than what is expected in the Thomson regime, where $Y_{\text{SSC}} \sim \sqrt{(\epsilon_e/\epsilon_B)}$ (Sari & Esin 2001), which means that the SSC cooling is suppressed in the Klein–Nishina regime when the photon energy measured in the electron rest frame is comparable to the electron rest mass. Thus, the SSC process is suppressed and enters into the deep KN regime given the parameters used in Figure 1. We conclude that it is challenging for the SSC scenario to achieve an energy flux level suggested by the MAGIC data. However, the situation is different if we consider the EIC scenario, where the external photons, e.g., coming from late-prompt emission, are more copious than synchrotron photons.

3. VHE γ -rays in the External Inverse-Compton Scenario

In this section, we present the EIC scenario, in which seed photons are attributed to extended and plateau emission. We fit the observed light curve of GRB 160821B using a phenomenological formula for extended emission (Kisaka et al. 2017),

$$L_{\text{EE}}(t) = L_{b,\text{EE}} \left(1 + \frac{t}{t_{\text{EE}}}\right)^{-\delta_{\text{EE}}}, \quad (1)$$

where $L_{b,\text{EE}} \simeq 6 \times 10^{48} \text{ erg s}^{-1}$ is the luminosity of the extended emission at the break energy $\epsilon_{b,\text{EE}}$, $t_{\text{EE}} \simeq 4 \times 10^2$ s is the duration of the extended emission, and $\delta_{\text{EE}} \simeq 10$ is the decay index that reflects the sharp decline of the extended emission (Troja et al. 2019). For plateau emission,

$$L_{\text{PL}}(t) = L_{b,\text{PL}} \left(\frac{t}{t_{\text{PL}}}\right)^{-\gamma_{\text{PL}}} \left(1 + \frac{t}{t_{\text{PL}}}\right)^{-\delta_{\text{PL}}}, \quad (2)$$

where $L_{b,\text{PL}} \simeq 4 \times 10^{43} \text{ erg s}^{-1}$ is the luminosity of the plateau emission at peak energy $\epsilon_{b,\text{PL}}$, $t_{\text{PL}} \simeq 2 \times 10^5$ s is the duration of the plateau emission, $\delta_{\text{PL}} = 20/3$ is the decay index to describe the decline of the plateau emission, and $\gamma_{\text{PL}} = 1/2$ is a factor to describe the gradual decline of the plateau emission. The index $\delta_{\text{EE/PL}}$ depends on the temporal evolution of the central engine activity. For example, the value of $\delta_{\text{EE/PL}} \sim 40/9$ can be derived assuming the fallback accretion rate, $\dot{M} \propto t^{-5/3}$ (Kisaka & Ioka 2015). We find that the X-ray light curve shows a sharp decline at the end of the extended and plateau emission, and $\delta_{\text{EE/PL}} > 40/9$ gives a better fit to the observation data, which could be realized by, e.g., the decay of magnetic fields via magnetic reconnections (Kisaka & Ioka 2015). The energy spectrum of late-prompt extended and plateau emission can generally be described by a broken power law $G(\epsilon) \propto (\epsilon/\epsilon_b)^{-\alpha+1}$ for $\epsilon < \epsilon_b$ and $G(\epsilon) \propto (\epsilon/\epsilon_b)^{-\beta+1}$ for $\epsilon > \epsilon_b$, where ϵ_b is the break energy, and α and β are the spectral indices that depend on the details of the emission mechanism (Kumar & Zhang 2014), such as the dissipative photosphere (Rees & Meszaros 2005), internal shock (Rees & Meszaros 1994), or magnetic dissipation (Zhang & Yan 2011). For simplicity, we assume that the late-prompt emission has an energy spectrum with a peak energy of $\epsilon_{b,\text{EE}} = 4 \times 10^3$ eV, $\alpha_{\text{EE}} = 0.5$, $\beta_{\text{EE}} = 2.6$, $\epsilon_{b,\text{PL}} = 30$ eV,

$\alpha_{\text{PL}} = 0.2$, and $\beta_{\text{PL}} = 2.7$ (Kagawa et al. 2019). Although the X-ray flux may overshoot the data around $10^{2.5}$ s, the results are not much affected because seed photons are mainly supplied by plateau emission thanks to the Klein–Nishina effect.

As shown in Figure 2, our EIC scenario can explain the $\sim 3\sigma$ detection by the MAGIC telescopes, considering observation uncertainties and poor weather conditions at earlier times. One can see that the SSC light curve is negligible compared to the EIC light curve (dashed–dotted line). The EIC light curve reaches a peak at the end of the extended emission because of the significant deceleration of the outflow at the deceleration time t_{dec} , where more nonthermal electrons are accumulated and IC scatterings cause time delays compared to seed photons impinging behind. Even though the EIC light curve is usually flatter than the SSC light curve (Murase et al. 2010, 2011), the transition of the seed photons coming from the extended emission to those from the plateau emission is seen in the EIC light curve. The predicted flux of VHE γ -rays can reach $\sim 10^{-12}$ erg cm $^{-2}$ s $^{-1}$ at 10^3 – 10^4 s. In Figure 2, we also show the corresponding Compton Y parameters, $Y_{\text{SSC}}(\gamma_e)$ and $Y_{\text{EIC}}(\gamma_e)$, as a function of electron energy. We find that the IC scattering with late-prompt plateau photons is between the Thomson regime and KN regime for electrons with Lorentz factors of $\gamma_e \sim 10^4$ – 10^6 . This is because the photon energy density $U_{\text{ph}}'[\varepsilon'_\gamma < \varepsilon'_{\text{KN}}]$ is insensitive to ε'_{KN} . Thus, we can expect EIC emission brighter than SSC emission.

The light curves in optical and near-infrared bands of GRB 160821B have a bump at ~ 1 day, which is interpreted as a kilonova/macronova component with a temperature of $T_{\text{KN,day}} \approx 4500$ K (Lamb et al. 2019; Troja et al. 2019). If the plateau emission is the internal origin and its dissipation region is inside the kilonova ejecta, electrons accelerated in the prolonged relativistic jets can emit high-energy gamma rays by upscattering of kilonova photons (Kimura et al. 2019). The thermal photon energy in the kilonova can be approximated to be $T_{\text{KN}}(t) \approx T_{\text{KN,day}}(t/\text{day})^{-0.8} \simeq 2.8 \times 10^4$ K at the time of MAGIC detection. Then, the KN break energy of the upscattered kilonova photon spectrum is estimated to be $\varepsilon_{\text{KN}} \approx m_e^2 c^4 / (2.8k_B T_{\text{KN}}) \sim 40$ GeV. This is well below the MAGIC threshold energy, and it is thus challenging to explain MAGIC data by the upscattered kilonova photons.

4. Summary and Discussion

We studied the origin of VHE γ -rays from GRB 160821B. While the SSC scenario has difficulty in explaining the MAGIC data, our results showed that the predicted EIC flux can dominate over the SSC flux in the case of GRB 160821B. In the EIC scenario, the extended and the plateau emission is attributed to the late-prompt emission from the long-lasting central engine, which can provide seed photons necessary for the EIC process. The resulting VHE γ -ray flux can reach $\sim 10^{-12}$ erg cm $^{-2}$ s $^{-1}$, which is consistent with the MAGIC observation.

In general, the detection of VHE γ -rays from short GRBs is more challenging, because they are less energetic than long GRBs. This is especially the case in the SSC scenario. This work demonstrated that the VHE signal can be enhanced by extended and plateau emission that can provide seed photons allowing the EIC emission to be dominant. One prominent feature of the EIC light curve that is accompanied by the extended and plateau emission is that the TeV peak is reached

around the end of the seed photons. Such EIC VHE γ -rays are promising targets for future IACTs observing short GRBs, and their off-axis emissions are promising EM counterparts of GWs from NS mergers (Murase et al. 2018). The detection of VHE γ -rays from short GRBs will also enable us not only to probe the activities of a central engine, which can be a black hole or NS, but also to constrain the environments of short GRBs and their host galaxies, as well as intergalactic magnetic fields in cosmic voids (Murase et al. 2009).

We thank Ke Fang, Susumu Inoue, and Koji Noda for useful discussion. The work of K.M. is supported by the Alfred P. Sloan Foundation, NSF grant No. AST-1908689, and KAKENHI No. 20H01901 and No. 20H05852. B.T.Z. acknowledges the IGC fellowship. C.C.Y. and P.M. acknowledge support from the Eberly Foundation. S.S.K. acknowledges the JSPS Research Fellowship, JSPS KAKENHI grant No. 19J00198. The early part of this work is supported by the Fermi GI program 111180.

ORCID iDs

B. Theodore Zhang  <https://orcid.org/0000-0003-2478-333X>
Kohta Murase  <https://orcid.org/0000-0002-5358-5642>
Chengchao Yuan  <https://orcid.org/0000-0003-0327-6136>
Shigeo S. Kimura  <https://orcid.org/0000-0003-2579-7266>
Peter Mészáros  <https://orcid.org/0000-0002-4132-1746>

References

- Abbott, B. P., Abbott, R., Abbott, T. D., et al. 2017a, *ApJL*, 848, L13
Abbott, B. P., Abbott, R., Abbott, T. D., et al. 2017b, *PhRvL*, 119, 161101
Abbott, B. P., Abbott, R., Abbott, T. D., et al. 2017c, *ApJL*, 848, L12
Abdalla, H., Adam, R., Aharonian, F., et al. 2019, *Natur*, 575, 464
Acciari, V. A., Ansoldi, S., Antonelli, L. A., et al. 2020, arXiv:2012.07193
Acciari, V. A., Ansoldi, S., Antonelli, L. A., et al. 2019a, *Natur*, 575, 455
Acciari, V. A., Ansoldi, S., Antonelli, L. A., et al. 2019b, *Natur*, 575, 459
Aleksić, J., Ansoldi, S., Antonelli, L. A., et al. 2016a, *Aph*, 72, 61
Aleksić, J., Ansoldi, S., Antonelli, L. A., et al. 2016b, *Aph*, 72, 76
Bartos, I., Corley, K. R., Gupte, N., et al. 2019, *MNRAS*, 490, 3476
Berger, E. 2014, *ARA&A*, 52, 43
Dai, Z.-G., Wang, X., Wu, X., & Zhang, B. 2006, *Sci*, 311, 1127
Derishev, E., & Piran, T. 2019, *ApJL*, 880, L27
Evans, P. A., Willingale, R., Osborne, J. P., et al. 2010, *A&A*, 519, A102
Fong, W.-f., Berger, E., Margutti, R., & Zauderer, B. A. 2015, *ApJ*, 815, L102
Ghisellini, G., Ghirlanda, G., Nava, L., & Firmani, C. 2007, *ApJL*, 658, L75
Gompertz, B., O'Brien, P., Wynn, G., & Rowlinson, A. 2013, *MNRAS*, 431, 1745
Gompertz, B. P., O'Brien, P. T., & Wynn, G. A. 2014, *MNRAS*, 438, 240
Hinton, J. 2009, *NJPh*, 11, 55005
Hinton, J., & Ruiz-Velasco, E. 2020, *JPhCS*, 1468, 012096
Inoue, S., Granot, J., O'Brien, P. T., et al. 2013, *Aph*, 43, 252
Inoue, S., Nava, L., Noda, K., et al. 2019, *Proc. ICRC*, 36, 703
Kagawa, Y., Yonetoku, D., Sawano, T., et al. 2019, *ApJ*, 877, 147
Kakuwa, J., Murase, K., Toma, K., et al. 2012, *MNRAS*, 425, 514
Kimura, S. S., Murase, K., Ioka, K., et al. 2019, *ApJL*, 887, L16
Kisaka, S., & Ioka, K. 2015, *ApJL*, 804, L16
Kisaka, S., Ioka, K., & Sakamoto, T. 2017, *ApJ*, 846, 142
Kneiske, T. M., Bretz, T., Mannheim, K., & Hartmann, D. H. 2004, *A&A*, 413, 807
Kumar, P., & Zhang, B. 2014, *PhR*, 561, 1
Lamb, G. P., Tanvir, N. R., Levan, A. J., et al. 2019, *ApJ*, 883, 48
Levan, A. J., Wiersema, K., Tanvir, N. R., et al. 2016, *GCN*, 19846, 1
Lorenz, E., Wagner, R., & Wagner, R. 2012, *EPJH*, 37, 459
Matsumoto, T., Kimura, S. S., Murase, K., & Mészáros, P. 2020, *MNRAS*, 493, 783
Mészáros, P. 2006, *RPPH*, 69, 2259
Mészáros, P., Fox, D. B., Hanna, C., & Murase, K. 2019, *NatRP*, 1, 585
Mészáros, P., & Rees, M. J. 1994, *MNRAS*, 269, L41
Mészáros, P., & Rees, M. J. 1997, *ApJ*, 476, 232
Murase, K., & Bartos, I. 2019, *ARNPS*, 69, 477

- Murase, K., Toma, K., Yamazaki, R., & Meszaros, P. 2011, [ApJ](#), 732, 77
- Murase, K., Toma, K., Yamazaki, R., Nagataki, S., & Ioka, K. 2010, [MNRAS](#), 402, L54
- Murase, K., Toomey, M. W., Fang, K., et al. 2018, [ApJ](#), 854, 60
- Murase, K., Zhang, B., Takahashi, K., & Nagataki, S. 2009, [MNRAS](#), 396, 1825
- Norris, J. P., & Bonnell, J. T. 2006, [ApJ](#), 643, 266
- Rees, M. J., & Meszaros, P. 1994, [ApJL](#), 430, L93
- Rees, M. J., & Meszaros, P. 2005, [ApJ](#), 628, 847
- Sari, R., & Esin, A. A. 2001, [ApJ](#), 548, 787
- Sari, R., Piran, T., & Narayan, R. 1998, [ApJL](#), 497, L17
- Schroeder, G., Margalit, B., Fong, W.-f., et al. 2020, [ApJ](#), 902, 82
- Troja, E., Castro-Tirado, A. J., Becerra González, J., et al. 2019, [MNRAS](#), 489, 2104
- Vanthieghem, A., Lemoine, M., Plotnikov, I., et al. 2020, [Galax](#), 8, 33
- Wang, X.-Y., Dai, Z., & Lu, T. 2001, [ApJ](#), 556, 1010
- Zhang, B., & Mészáros, P. 2001, [ApJ](#), 559, 110
- Zhang, B., & Yan, H. 2011, [ApJ](#), 726, 90
- Zhang, B. T., Murase, K., Veres, P., & Mészáros, P. 2020, arXiv:2012.07796